Improved fitting of periodic variable star light curves through regularized regression

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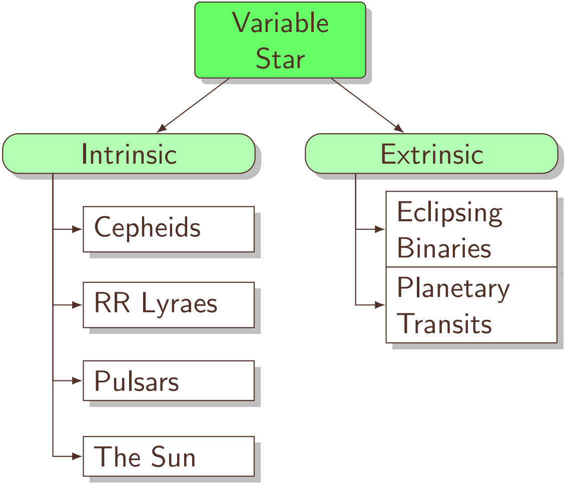
ASNY 2015 – November 7th, 2015

Variable Stars

Overview

- in general, any star whose brightness changes on short timescales is a variable star
- many different types exist

Some Classes of Variable Stars



Pulsating Periodic Intrinsic Variables

For the remainder of this talk:

variable star $\equiv$ pulsating periodic intrinsic variable star.

- not in hydrostatic equilibrium
  - typically in the instability strip
- periodic oscillation
  - predictable
- stellar pulsation
  - κ-mechanism

Henrietta Swan Leavitt

Henrietta Swan Leavitt

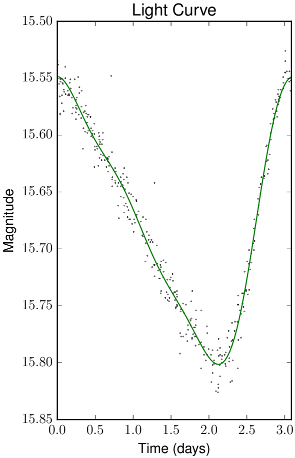
- worked as a “computer” at Harvard in the early 20th century
- discovered a relation between the period and luminosity of Cepheids
  - Leavitt’s law
  - standard candles
- enabled Edwin Hubble to discover the expansion of the Universe

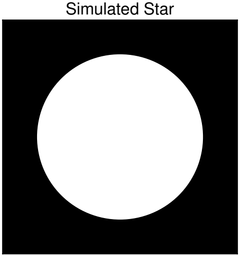
Light Curves

Overview

- repeated photometric measurements of an object over time
- plotting brightness versus time gives us a light curve

Light Curve of a Cepheid Variable Star





Visualization of OGLE-LMC-CEP-0002

Fourier Analysis

Joseph Fourier

- any continuous, periodic function can be represented as an infinite Fourier series

$$f(t) = A_0 + \sum_{k=1}^{\infty} A_k \cos(k\omega t + \Phi_k)$$

- characterized by the angular frequency ω, the mean A_0, the amplitudes A_k, and the phase shifts Φ_k

Fourier Analysis of Periodic Light Curves

$$m(t) = A_0 + \sum_{k=1}^{n} A_k \cos(k\omega t + \Phi_k)$$

- Cepheid-like light curves well described by nth order Fourier Series
- physically they are close to harmonic oscillators

Solving for Series Parameters

$$m(t) = A_0 + \sum_{k=1}^{n} A_k \cos(k\omega t + \Phi_k)$$

- Fourier series are non-linear
  - simultaneously finding the optimal n, ω, A_k, and Φ_k is not easy
- we must break the problem into easier sub-problems

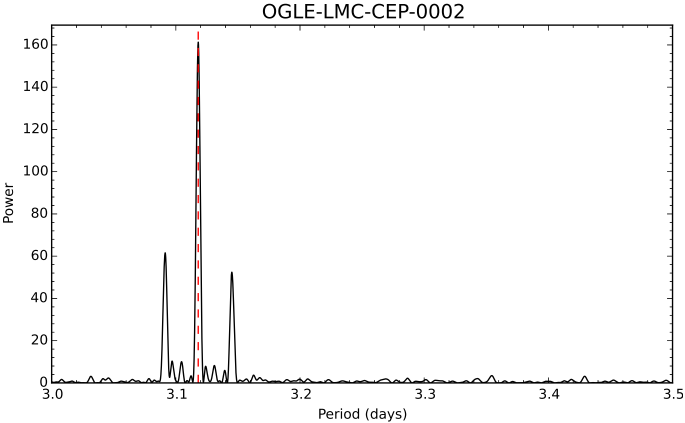
Period finding

- the most important parameter is the period

$$\omega = 2\pi/P$$

- we can approximate this by itself using a periodogram
  - Lomb-Scargle

Lomb-Scargle Periodogram



Periodogram of star with 3.11804 day period

Linearizing Phase Shift

$$m(t) = A_0 + \sum_{k=1}^{n} A_k \cos(k\omega t + \Phi_k)$$

- Φ_k still makes this a non-linear optimization problem
- trig identities to the rescue!

Linearizing Phase Shift (continued)

$$\cos(\alpha \pm \beta) = \cos(\alpha)\cos(\beta) \mp \sin(\alpha)\sin(\beta)$$

$$\begin{aligned} A_k \cos(k\omega t + \Phi_k) &= A_k \cos(\Phi_k)\cos(k\omega t) - A_k \sin(\Phi_k)\sin(k\omega t) \\ &= a_k \sin(k\omega t) + b_k \cos(k\omega t) \end{aligned}$$

It's Linear!

$$m(t) = A_0 + \sum_{k=1}^{n} \left[a_k \sin(k\omega t) + b_k \cos(k\omega t) \right]$$

can be written in the form

$$\mathbf{X}\vec{\beta} = \vec{y}$$

which can be approximated using ordinary linear regression

System of Equations

$$\vec{y} \to \begin{pmatrix} m_1 & m_2 & \dots & m_N \end{pmatrix}$$

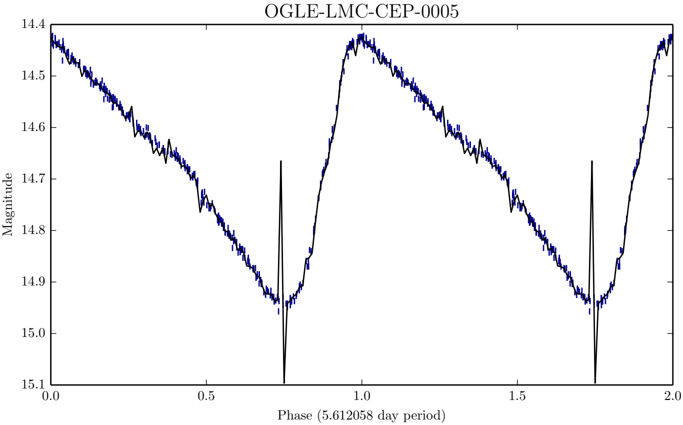
$$\vec{\beta} \to \begin{pmatrix} A_0 & a_1 & b_1 & \dots & a_n & b_n \end{pmatrix}$$

$$\mathbf{X} \to \begin{pmatrix} 1 & \sin(1\omega t_1) & \cos(1\omega t_1) & \dots & \sin(n\omega t_1) & \cos(n\omega t_1) \\ \vdots & \vdots & \vdots & \ddots & \vdots & \vdots \\ 1 & \sin(1\omega t_N) & \cos(1\omega t_N) & \dots & \sin(n\omega t_N) & \cos(n\omega t_N) \end{pmatrix}$$

How many terms?

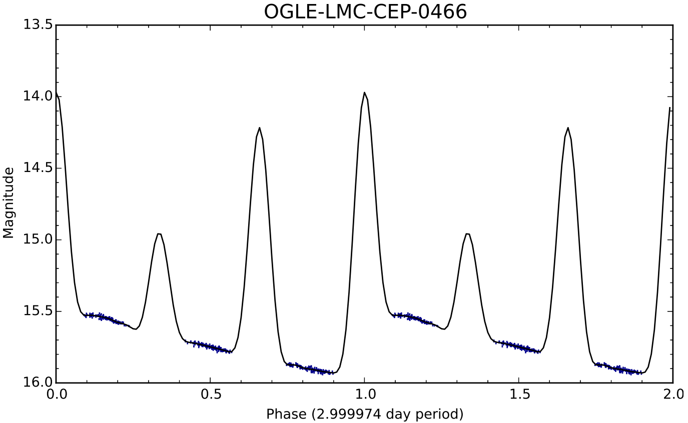
- wait, we never decided on the order of the fit, n
- it's just a truncated series expansion
    - more terms means better, right?
    - let's try 100 terms…

Overfitting



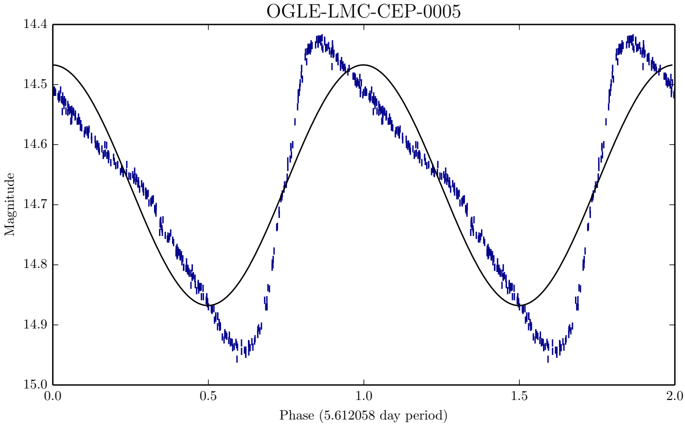
100th order fit

Overfitting (again)



12th order fit

Underfitting



1st order fit

Choosing n

- need some criteria to decide the order of the fit
- Baart's criteria is often used for this
  - iterative approach, increasing n until diminishing returns
  - good at avoiding underfitting
  - bad at avoiding overfitting

Taking a step back

- take photometric measurements
- find the period
- linearize
- approximate coefficients with OLS
- find the best order of fit using Baart's criteria

Taking a step back

- take photometric measurements
- **periodogram**
- linearize
- **regression**
- **model selection**

Plotypus

- tool for modeling and plotting light curves
- free and open source
- version controlled and documented
- generated the light curve plots in this presentation
- astroswego.github.io/plotypus/
- download today!

Earl Bellinger, Daniel Wysocki, Shashi Kanbur, 2015–

Unconstrained Regression

$$\mathbf{X}\vec{\beta} = \vec{y}$$

$$\begin{aligned}(A_0, a_k, b_k) &= \operatorname*{argmin}_{\beta} \left\| \mathbf{X}\vec{\beta} - \vec{y} \right\|_2^2 \\ &= \operatorname*{argmin}_{(A_0, a_k, b_k)} \sum_{i=1}^{N} \left(A_0 + \sum_{k=1}^{n} \left[\begin{array}{r} a_k \sin(k\omega t_i) \\ +b_k \cos(k\omega t_i) \end{array} \right] - m_i \right)^2\end{aligned}$$

Find the coefficients which minimize the residual sum of squares

ℓ_0 Regularization

$$(A_0, a_k, b_k) = \underset{\beta}{\operatorname{argmin}} \left\{ \left\| \mathbf{X}\vec{\beta} - \vec{y} \right\|_2^2 + \lambda \left\| \vec{\beta} \right\|_0 \right\}$$

- $\left\| \vec{\beta} \right\|_0$ is equal to the number of non-zero terms in $\vec{\beta}$
- adds a penalty on the number of parameters, weighted by λ
- this is computationally expensive

ℓ_1 Regularization (LASSO)

$$(A_0, a_k, b_k) = \underset{\beta}{\operatorname{argmin}} \left\{ \left\| \mathbf{X}\vec{\beta} - \vec{y} \right\|_2^2 + \left\| \vec{\beta} \right\|_1 \right\}$$

$$= \underset{(A_0, a_k, b_k)}{\operatorname{argmin}} \left\{ \begin{array}{l} \displaystyle\sum_{i=1}^{N} \left(A_0 + \sum_{k=1}^{n} \left[\begin{array}{r} a_k \sin(k\omega t_i) \\ + b_k \cos(k\omega t_i) \end{array} \right] - m_i \right)^2 \\ \displaystyle + \lambda \sum_{k=0}^{n} |A_k| \end{array} \right\}$$

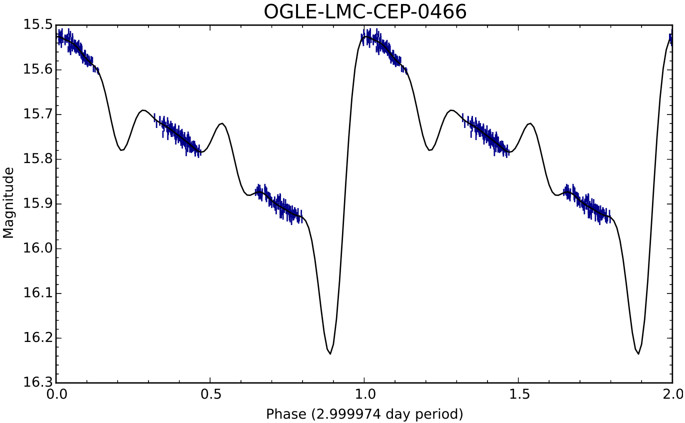
- least absolute shrinkage and selection operator (LASSO)
- adds a penalty on the sum of the amplitudes, weighted by λ
- automatically zeroes out non-contributing terms

Model Selection with Grid Search

- use grid search with cross-validation
  - search over the order of fit n
- cross-validation helps fit underlying function, not just the data

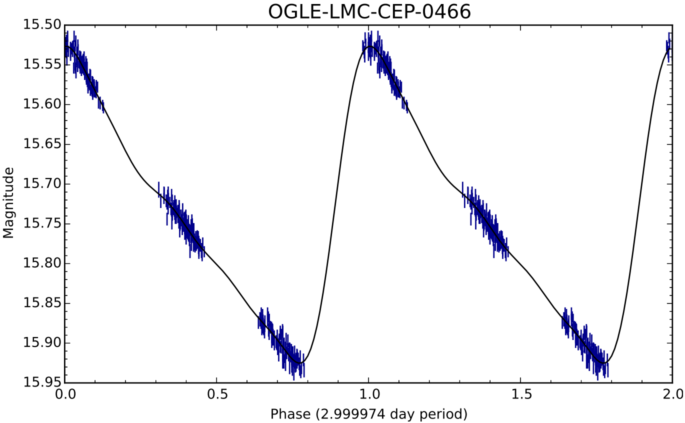
Results

OLS/Baart Light Curve



10th order fit using OLS/Baart.

Lasso Light Curve



10th order fit using Lasso/Grid Search.

Performance of LASSO/Grid Search versus OLS/Baart

Galaxy	Type	Stars	N (SD)	LASSO R^2 (MAD)	Baart R^2 (MAD)	Significance
(all)	(all)	52844	643.1 (462.0)	**0.8594 (0.1741)**	0.8492 (0.1864)	$p < .0001$
(all)	CEP	7999	740.1 (298.4)	**0.9816 (0.0191)**	0.9810 (0.0198)	$p < .0001$
(all)	T2CEP	596	747.6 (612.0)	**0.9145 (0.1159)**	0.9009 (0.1328)	$p < .0001$
(all)	ACEP	89	497.3 (225.0)	**0.9700 (0.0245)**	0.9689 (0.0267)	$p < .0001$
(all)	RRLYR	44160	624.4 (481.6)	**0.8316 (0.1816)**	0.8197 (0.1926)	$p < .0001$
LMC	(all)	28491	522.3 (227.7)	**0.7812 (0.1695)**	0.7723 (0.1779)	$p < .0001$
LMC	CEP	3342	536.8 (219.7)	**0.9840 (0.0172)**	0.9833 (0.0180)	$p < .0001$
LMC	T2CEP	201	538.3 (232.6)	**0.8672 (0.1569)**	0.8599 (0.1653)	$p < .0001$
LMC	ACEP	83	477.3 (214.7)	**0.9704 (0.0233)**	0.9701 (0.0245)	$p < .0001$
LMC	RRLYR	24865	520.3 (228.6)	**0.7544 (0.1667)**	0.7452 (0.1755)	$p < .0001$
SMC	(all)	7146	851.4 (256.7)	**0.9109 (0.1241)**	0.9091 (0.1266)	$p < .0001$
SMC	CEP	4625	886.5 (256.2)	**0.9800 (0.0195)**	0.9796 (0.0200)	$p < .0001$
SMC	T2CEP	42	891.2 (241.4)	**0.7965 (0.2235)**	0.7888 (0.2379)	$p < .0001$
SMC	ACEP	6	774.3 (190.2)	0.9277 (0.0709)	0.9272 (0.0706)	$p = 0.2188$
SMC	RRLYR	2473	785.2 (244.8)	**0.6299 (0.1915)**	0.6203 (0.1962)	$p < .0001$
BLG	(all)	17207	756.8 (698.1)	**0.9579 (0.0445)**	0.9527 (0.0514)	$p < .0001$
BLG	CEP	32	824.2 (569.0)	**0.9742 (0.0342)**	0.9703 (0.0396)	$p < .0001$
BLG	T2CEP	353	849.7 (746.8)	**0.9525 (0.0643)**	0.9457 (0.0747)	$p < .0001$
BLG	RRLYR	16822	754.7 (697.2)	**0.9581 (0.0440)**	0.9528 (0.0509)	$p < .0001$

Median coefficients of determination (R^2) and median absolute deviations (MAD) for models selected by cross-validated LASSO and Baart's ordinary least squares on OGLE I-band photometry. P-values obtained by paired Mann-Whitney U tests.

Missing Harmonics

- LASSO makes no distinction between higher and lower order terms
  - if it doesn't contribute, it goes to zero
- this can result in $A_i = 0$, when $A_j \neq 0$, $j > i$
  - contrary to pulsation models, which say amplitude decreases with order

$$A_1 > A_2 > \ldots > A_n$$

- explanations:
  - harmonics absent from observations
    - e.g. we observe only near zero-crossing
  - interference pattern in pulsation (gets political)
  - others? (please tell me)

Multifrequency Variable Stars

$$m(t) = A_0 + \sum_{k_1=-n}^{n} \dots \sum_{k_p=-n}^{n} A_{\mathbf{k}} \cos\big((\mathbf{k} \cdot \boldsymbol{\omega})t + \Phi_{\mathbf{k}}\big)$$

$$\mathbf{k} \to \begin{pmatrix} k_1 & \dots & k_p \end{pmatrix} \quad \boldsymbol{\omega} \to \begin{pmatrix} \omega_1 & \dots & \omega_p \end{pmatrix}$$

- some variable stars oscillate with multiple (p) periods
- OLS fails to accurately fit these light curves
  - tools exist to manually fix certain amplitudes to zero
- LASSO successful in automatically zeroing out amplitudes

Questions?